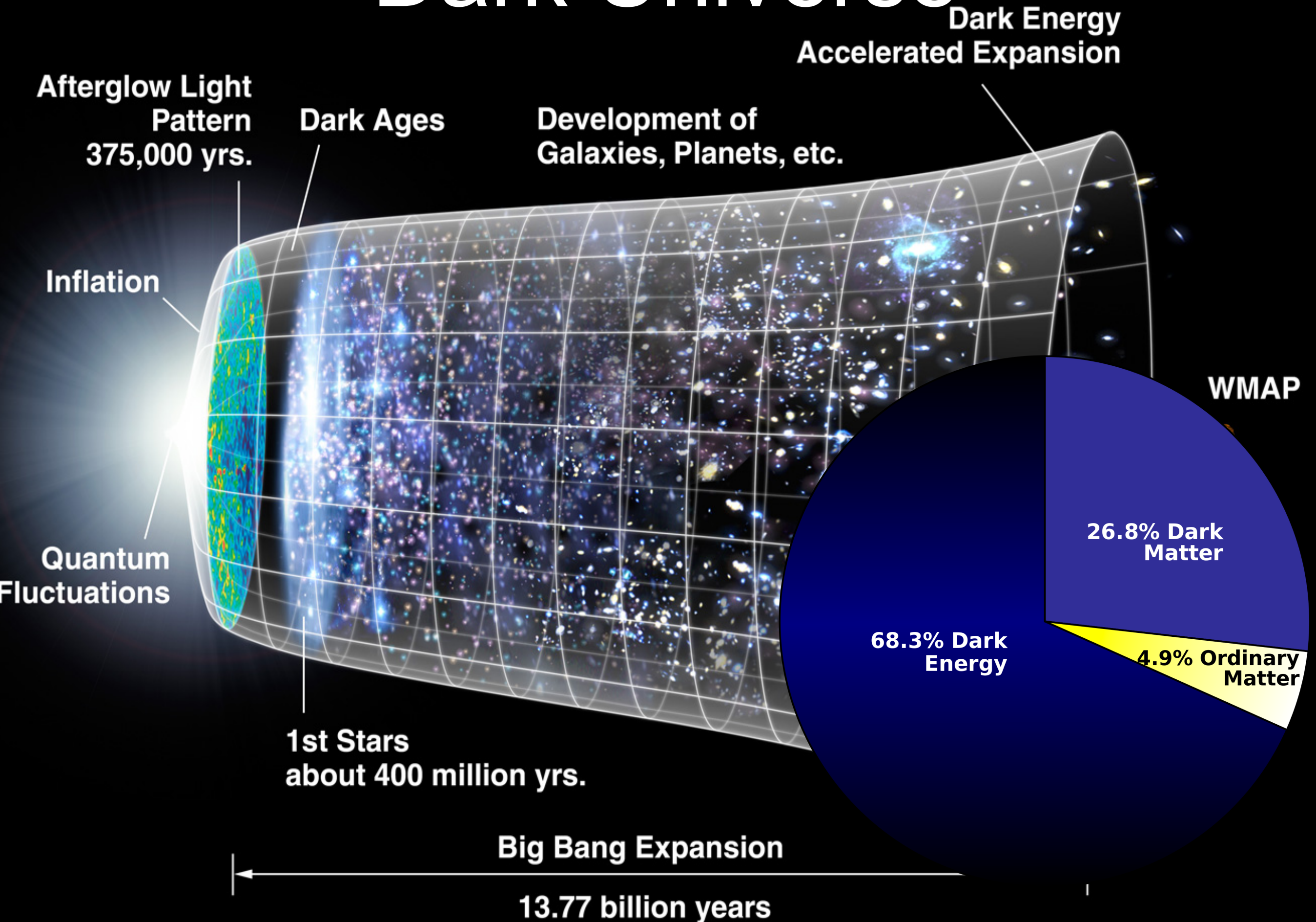


Pseudo phase space density of dark matter halos in the dwarf spheroidals

**Kohei Hayashi (JSPS fellow)
ICRR, The University of Tokyo**

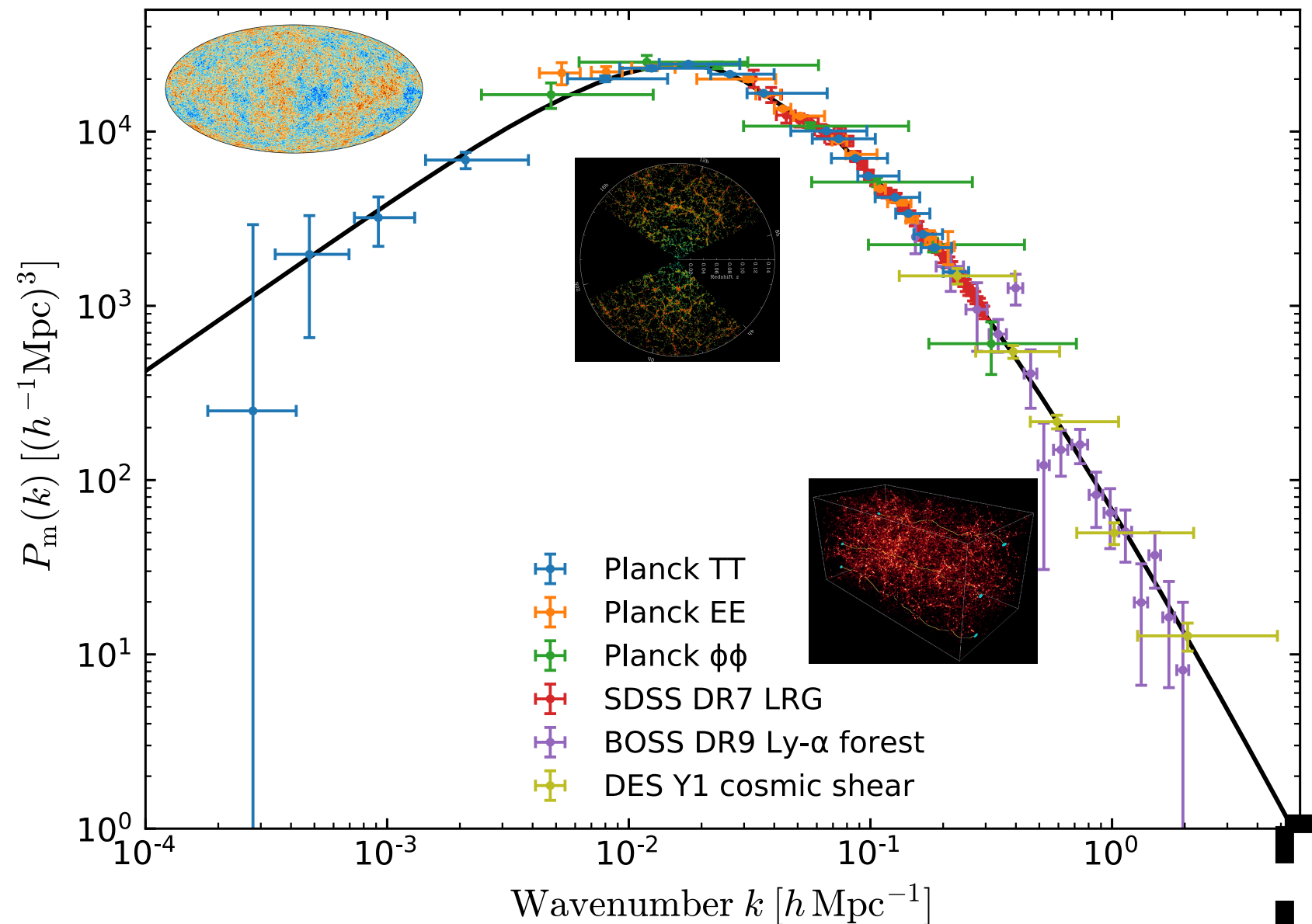
Collaborators: Tomoaki Ishiyama (Chiba Univ.), Keiichi Umetsu (ASIAA),
Masashi Chiba (Tohoku Univ.)

Dark Universe



Λ Cold Dark Matter model

ESA and the Planck Collaboration



Large scale ($> 1 \text{Mpc}$)
 $\Rightarrow$ Remarkable success!

Small scale ($< 1 \text{Mpc}$)
 $\Rightarrow$ What's going on?

?

Dark matter halo structure

THE STRUCTURE OF COLD DARK MATTER HALOS

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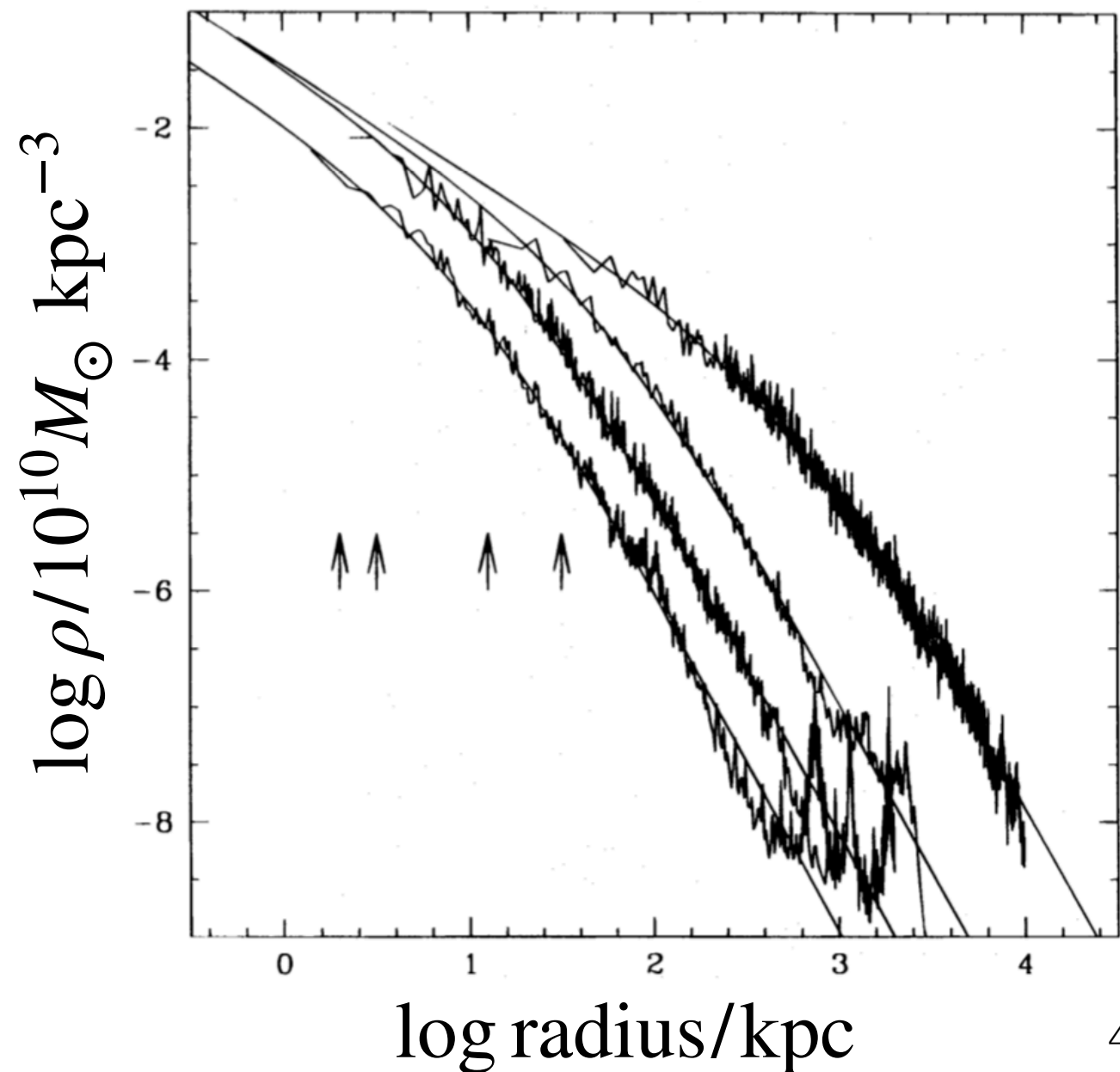
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AND

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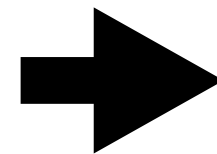
Received 1995 August 1; accepted 1995 December 4



NFW profile:

$$\frac{\rho(r)}{\rho_{\text{crit}}} = \frac{\delta_c}{(r/r_s)(1 + r/r_s)^2}$$

$$\delta_c = \frac{200}{3} \frac{c^3}{[\ln(1 + c) - c/(1 + c)]}$$



Universal profile

Pseudo phase space density (PPSD) profile

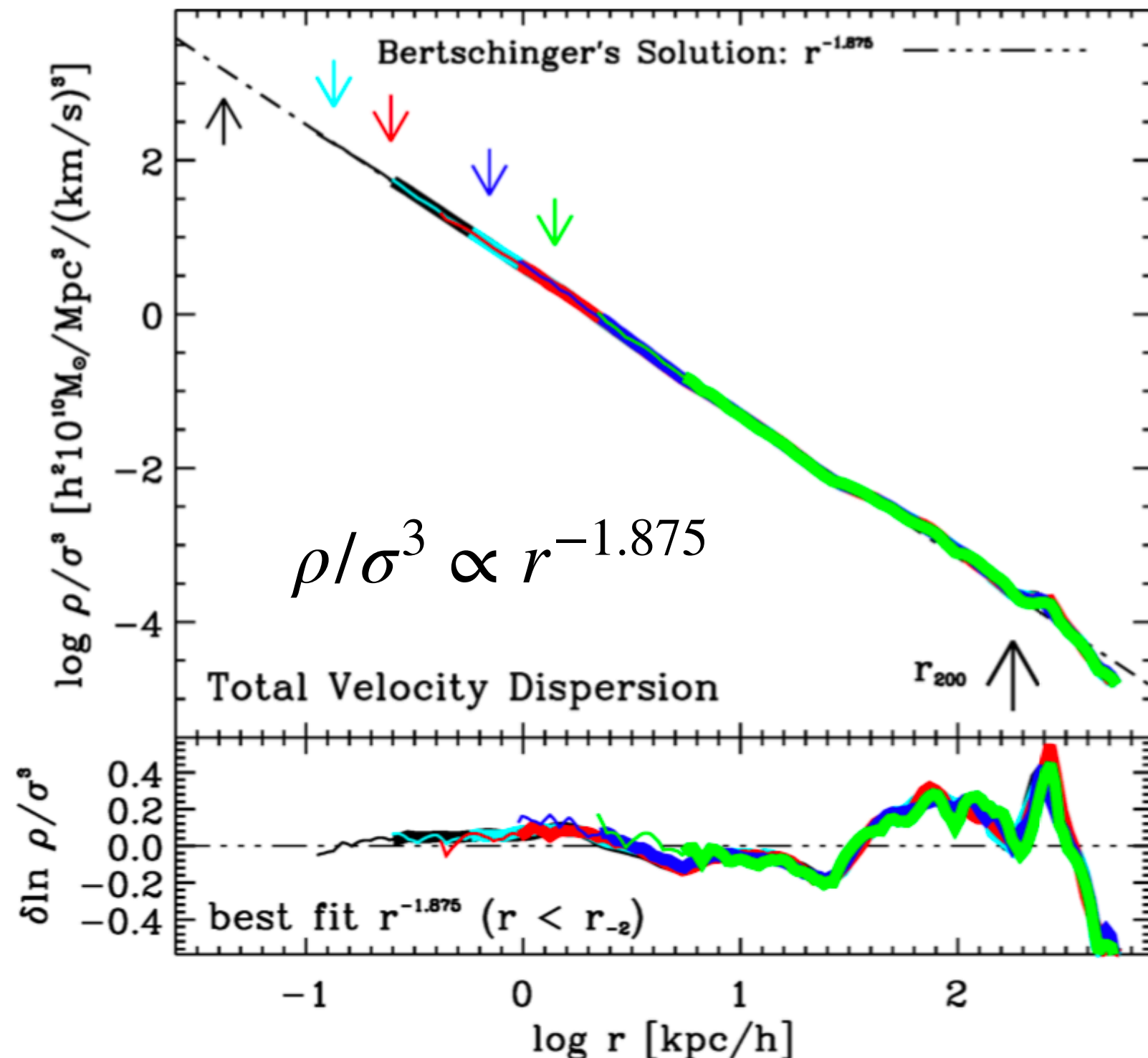
Navarro et al. (2010)

1. Universal profile (Taylor & Navarro 2001)

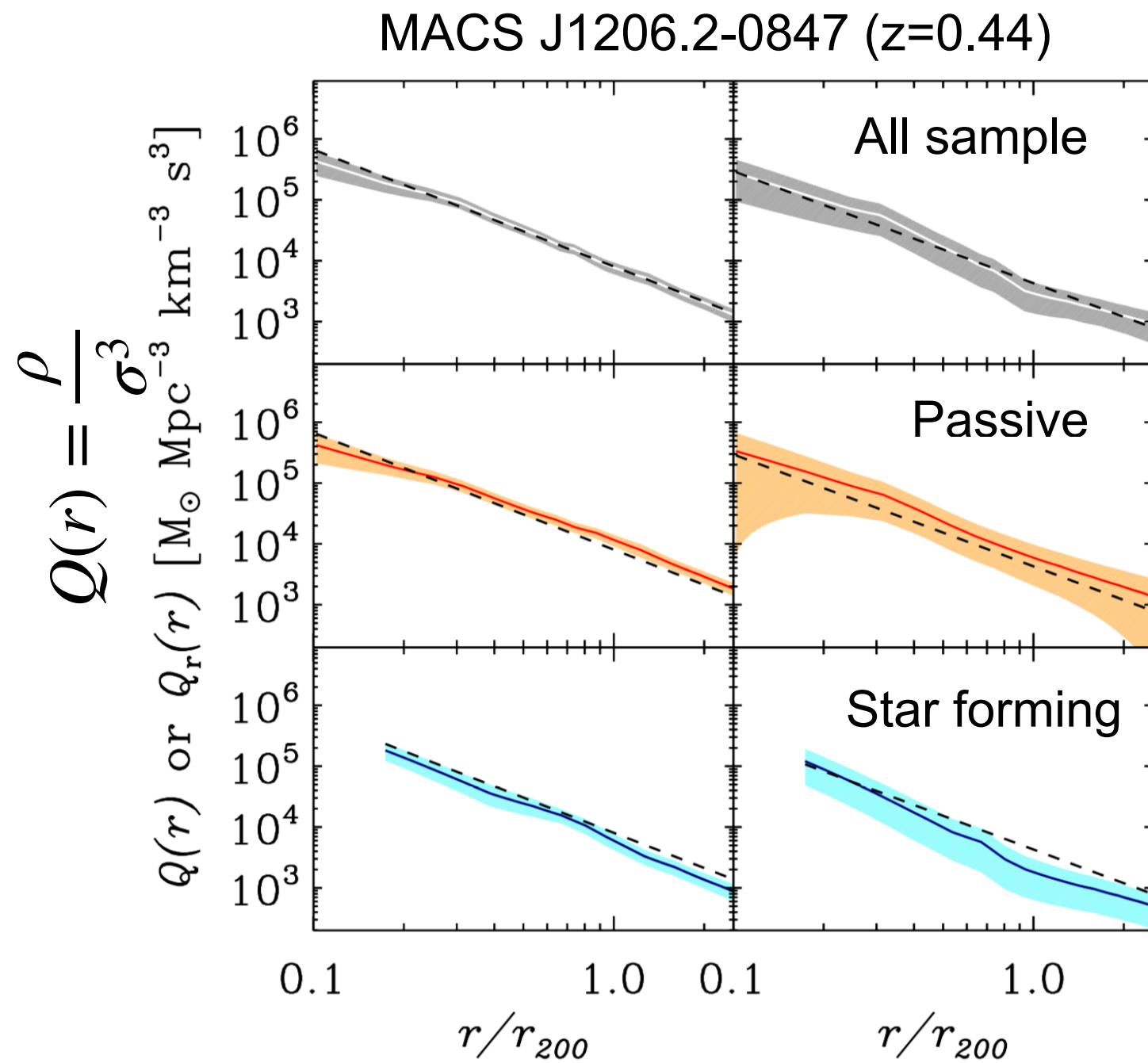
$$\rho/\sigma^3 \propto r^{-1.875}$$

2. The power-law is not largely dependent upon “cosmology” and slopes of dark matter density profiles (e.g., Navarro+09; Ludlow+12).

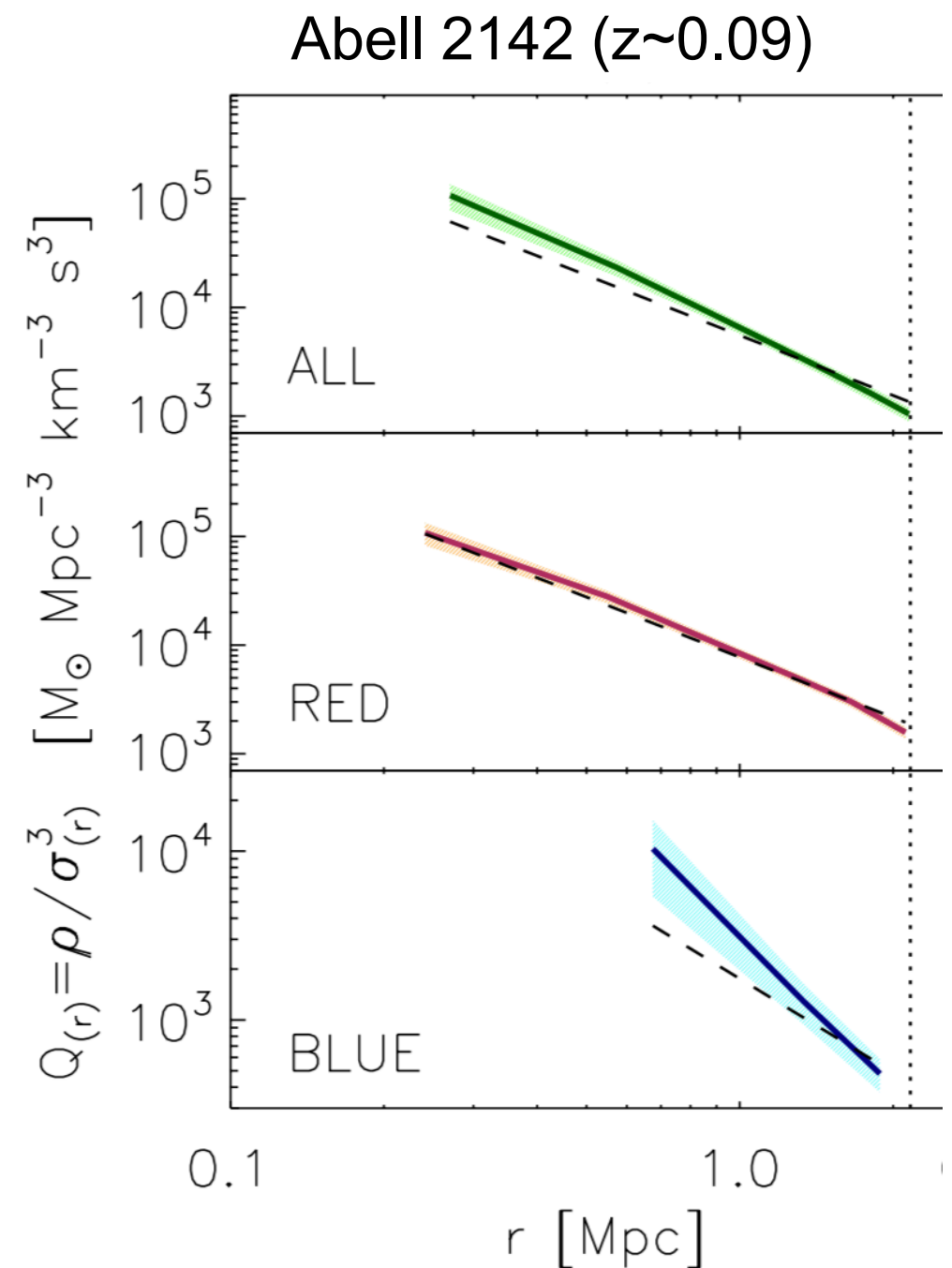
3. Identical to self-similar secondary infall solution in an Einstein-de Sitter universe (Fillmore & Goldreich 84; Bertschinger 85).



PPSD profile for galaxy cluster



Biviano+ (2013)



Munari+ (2014)

PPSD profile for dwarf galaxy?

Non-spherical dynamical mass models

Unobservable

Non-spherical dark matter density profile

$$\rho_{\text{DM}}(r) = \frac{\rho_s}{(r/r_s)^\gamma [1 + (r/r_s)^\alpha]^{(\beta-\gamma)/\alpha}}$$

$$r^2 = R^2 + \frac{z^2}{Q^2} \quad \leftarrow \text{DM halo axial ratio}$$

Non-spherical stellar profile

$$\rho_*(r_*) = \frac{3L}{4\pi r_p^3} \left[1 + \frac{r_*^2}{r_p^2} \right]^{-5/2}$$

$$r_*^2 = R^2 + \frac{z^2}{q^2} \quad \leftarrow \text{stellar axial ratio}$$

Axisymmetric Jeans equations

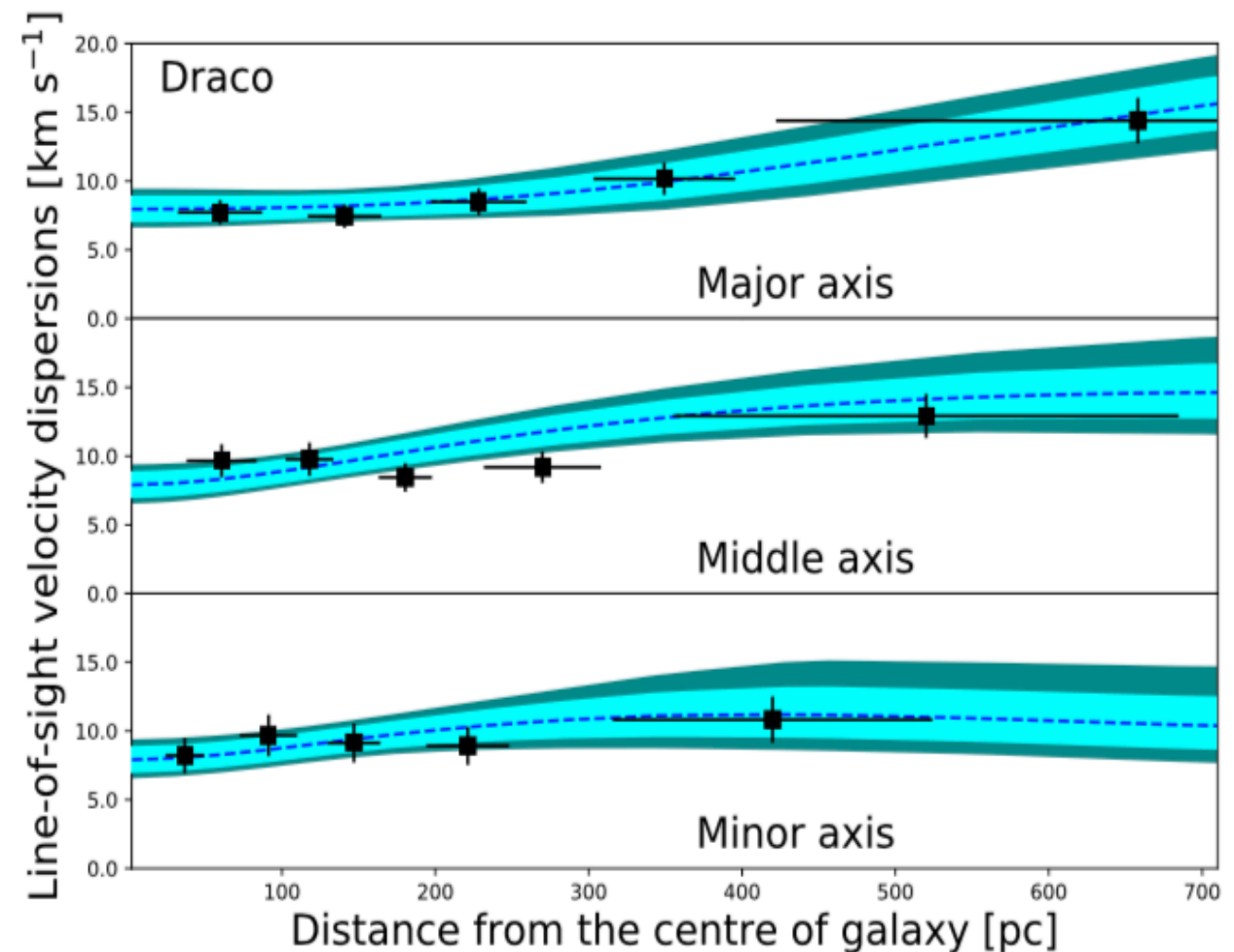
➔ $(\overline{v_z^2}, \overline{v_\phi^2})$ $\overline{v_R^2}$ is unknown parameter as $\beta_z = 1 - \overline{v_z^2}/\overline{v_R^2}$

$\sigma_{\text{l.o.s}}(\text{Theory})$

FIT

$\sigma_{\text{l.o.s}}(\text{obsrved})$

Draco dSph

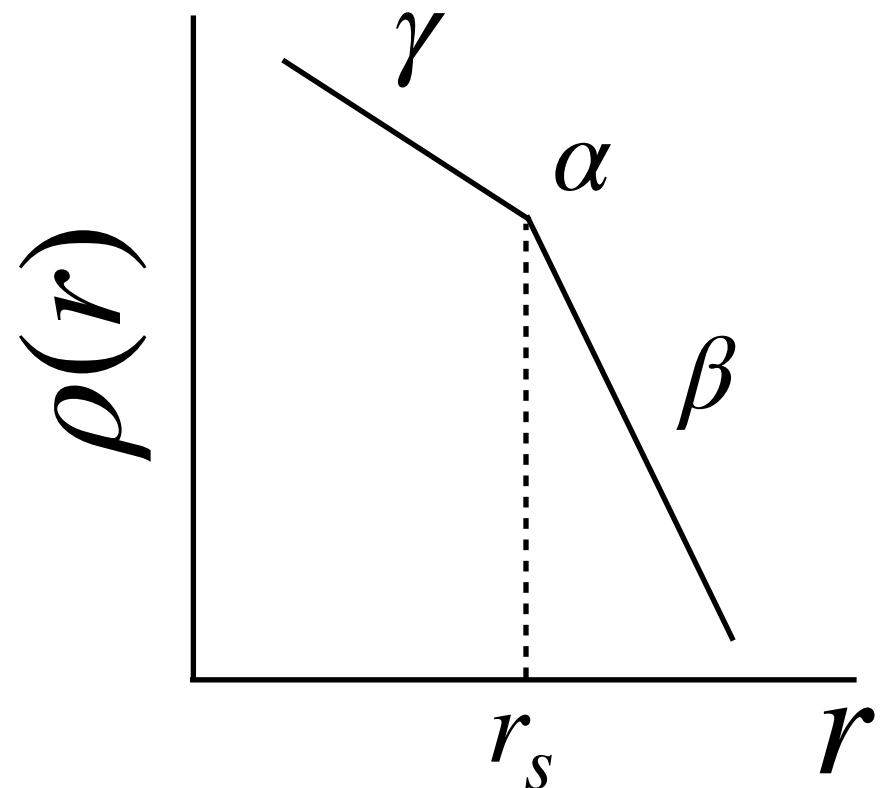


DM profiles of the classical dwarfs

- Parameterized dark matter density profile

$$\rho_{\text{DM}}(r) = \frac{\rho_s}{(r/r_s)^\gamma [1 + (r/r_s)^\alpha]^{(\beta-\gamma)/\alpha}}$$

$$r^2 = R^2 + \frac{z^2}{Q^2}$$



DM profiles of the classical dwarfs

KH+ (in prep.)

Preliminary

Cuspy profile

Sorry...

Less cuspy profile

Sorry...

Dark matter density [$M_{\text{sun}}/\text{kpc}^3$]

Major axis [kpc]

PPSD profiles for dSphs

KH+ (in prep.)

$$Q(r) = \frac{\rho}{\sigma^3} \quad \sigma^3 = \sigma_R \sigma_z \sigma_\phi$$

Preliminary

Sorry...

Pseudo phase space density [Msun/kpc³/(km/s)³]

Major axis [kpc]

Draco dSphs

Sorry...

What's an origin of this deviation of PPSD profile?

What's an origin of this deviation of PPSD profile?

- **Systematics?**

Velocity dispersion is derived by not dark matters but star kinematics.

- **Baryonic feedback?**

but, Draco prefers to have cusped DM halo. SF feedback might be not efficient...

- **Alternative DM models?**

(WDM, Axion, SIDM,...)

- still ongoing discussion ...

Sorry...

Summary

- Pseudo phase space density profile is universal within the framework of Λ CDM universe.
- The slope of PPSD profile is identical to self-similar secondary infall solution in an Einstein-de Sitter universe.
- We look into PPSD profiles on dwarf-galaxy scales for the first time.
- We find **the diversity of inner slopes of DM density profiles** in the classical dSphs and **their PPSD profiles deviate from Λ CDM models.**